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Temperature; 3.3.2 The Escape of Gas Heated by Ionization; 3.3.3 The Escape of Gas Following a Supernova Explosion; 3.3.4 Population II.5; 3.4 The First Star Clusters; 3.4.1 Clusters of Population III Stars and of Metal-Poor Stars; 3.4.2 The Origin of Globular Clusters; 3.5 The First Galaxies; 3.6 The First Active Galactic Nuclei; 3.6.1 Population III Black Holes; 3.6.2 Black-Hole Mergers; 3.6.3 The Highest-Redshift QSOs; 3.6.4 Direct Collapse to Black Holes 3.7 Low-Metallicity HII Regions 3.8 Numerical Techniques and Their Limitations; 3.8.1 Collisionless Dynamics; 3.8.2 Collisionless Dynamics: Particle-Mesh Codes; 3.8.3 Collisionless Dynamics: Treecodes; 3.8.4 Gas Dynamics; 3.8.5 Gas Dynamics: Smooth Particle Hydrodynamics; 3.8.6 Gas Dynamics: Eulerian Codes; 3.8.7 Improving Resolution Through Mesh Refinement; 3.8.8 Radiative Transfer; 3.9 Hints for Further Study; 4 Cosmic Reionization; 4.1 Overview; 4.2 The Properties of the Sources of Reionization; 4.2.1 The Surface Brightness of Reionization Sources; 4.2.2 Reionization in a Hydrogen-Only IGM 4.2.3 Reionization in a Hydrogen-Helium IGM 4.2.4 Results for a Homogeneous IGM; 4.2.5 Mean Metallicity at Reionization; 4.3 Adding Realism to the Calculations; 4.3.1 Escape of Ionizing Photons; 4.3.2 Clumpy IGM; 4.3.3 Two-Parameter Models; 4.4 Luminosity Function of Ionizing Sources; 4.4.1 Detecting Lyman α from Ionizing Sources; 4.5 Reionization by Population III Stars; 4.6 How Is the Intergalactic Medium Enriched?; 4.7 Reheating of the Intergalactic Medium; 4.8 Keeping the Intergalactic Medium Ionized; 4.9 Hints for Further Study; Part 2 Observational Techniques and their Results 5 Studying the Epoch of Reionization of Hydrogen

Sommario/riassunto

This up-to-date and concise account of a critical period of the early universe directly links the latest theories and experiments. Targeted at cosmological problems rather than specific methods, it begins with an introduction reviewing the early universe and looks at why reionization is important. The process of reionization analyzes simple analytical considerations and compares existing observations, while a further chapter describes some of the issues regarding the transition from Population III to Population II stars, as well as the constraints that can be derived from WMAP. Further chapter

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